

Surface structure of the CoRoT^{*} CP2 target star HD 50773

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ABSTRACT

Aims. We compare surface maps of the chemically peculiar star HD 50773 produced with a Bayesian technique and based on high quality CoRoT photometry with those derived from rotation phase resolved spectropolarimetry. The goal is to investigate the correlation of surface brightness with surface chemical abundance distribution and the stellar magnetic surface field.

Methods. The rotational period of the star was determined from a nearly 60 day long continuous light curve obtained during the initial run of CoRoT. Using a Bayesian approach to star-spot modelling, which in this work is applied for the first time for the photometric mapping of a CP star, we derived longitudes, latitudes and radii of four different spot areas. Additional parameters like stellar inclination and the spot's intensities were also determined. The CoRoT observations triggered an extensive ground-based spectroscopic and spectropolarimetric observing campaign and enabled us to obtain 19 different high resolution spectra in Stokes parameters I and V with NARVAL, ESPaDOnS, and SemelPol spectropolarimeters. Doppler and Magnetic Doppler imaging techniques allowed us to derive the magnetic field geometry of the star and the surface abundance distributions of Mg, Si, Ca, Ti, Cr, Fe, Ni, Y, and Cu.

Results. We find a dominant dipolar structure of the surface magnetic field. The CoRoT light curve variations and abundances of most elements mapped are correlated with the aforementioned geometry: Cr, Fe, and Si are enhanced around the magnetic poles and coincide with the bright regions on the surface of HD 50773 as predicted by our light curve synthesis and confirmed by photometric imaging.

Key words. stars: atmospheres – stars: chemically peculiar – stars: individual: HD 50773 – stars: magnetic fields – stars: surface abundance structure

1. Introduction

CoRoT (Convection, Rotation and planetary Transits) is a space mission with the participation of ESA's Science Program and Research and Scientific Support Department (RSSD), Austria, Belgium, Brazil, Germany, and Spain. It focuses on high precision photometry from space, also taking advantage of observing given targets continuously during nearly half a year, which is impossible from the ground. A technical overview is presented in Boisnard & Auvergne (2006), and the asteroseismology related mission aspects are discussed in Baglin et al. (2006) and references therein.

HD 50773 (BD -00 1488, TYC 4801-2-1, mag(B) = 9.50) was observed during the initial run of CoRoT only little more than a month after the launch on December 27, 2006, on board of a Soyuz Fregat II-1b. The star was chosen in the seismology field as one of ten possible targets because of its location in the classical instability strip and of its suspected chemical peculiar-

ity. The reduction of the CoRoT photometry to the N2 data format is described in Appourchaux et al. (2008), and our additional reduction steps are given in Sect. 2 of this paper.

Not much has been published about this star, which is classified in SIMBAD as an A2 star and as an A4 – A9 suspected chemically peculiar star in Renson et al. (1991) with so far no measured magnetic field nor rotation period.

Magnetic Ap stars, to which HD 50773 belongs, represent about 1 % to 5 % of the upper main sequence stars and exhibit highly ordered, very stable and often very strong magnetic fields. They frequently show both brightness- and spectral line profile variations synchronised to stellar rotation. These variations are hitherto most successfully explained by the oblique rotator model (introduced by Stibbs 1950) and are attributed to oblique magnetic and rotation axes and to the presence of a non-uniform distribution of chemical elements on their surface.

The abundance inhomogeneities in turn are said to arise from the selective diffusion of ions under the competitive action of radiative acceleration and gravitational settling (Michaud 1970) under the influence of the (oblique) magnetic field, possibly in combination with a weak stellar wind (see, e.g., Babel 1992).

The light curve of HD 50773 obtained by CoRoT has a double wave form with two clear maxima of slightly different am-

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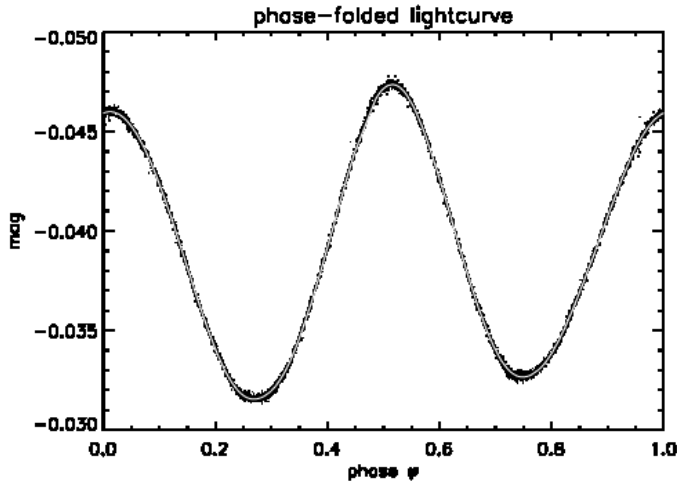


Fig. 1. Phase plot of HD 50773 using a binned light curve after the subtraction of effects due to the CoRoT orbit and the stellar signal not attributed to rotation. Overplotted (white line) is the four-spot model fit described in Sect. 4 of this Paper.

plitudes at phases of $\phi \approx 0.05$ and 0.52 (see Fig. 1). This photometric variability is, as mentioned, likely to be connected with the inhomogeneous surface element distribution as shown in e.g., Krtićka et al. (2007) and references therein. As the star seems to be of the Cr CP2-type, we expect brighter spots rather than dark ones (see e.g. Mikulášek et al. 2007 and Mikulášek et al. 2008). Bright photometric spots seem to be closely connected with overabundance regions of Si or Fe, and the mechanism of the origin of their contrast in the optical spectrum region is well described in Krtićka et al. (2009).

It is still not understood why some of the upper main sequence stars are magnetic and chemically peculiar and others are not, where the magnetic fields of CP2 stars originate from, and how these fields exactly interplay and correlate with the inhomogeneous surface distribution of chemical elements. Deriving in detail the surface structure of HD 50773 in the way presented in this paper will contribute to solving this puzzle.

Our paper is organised in Sect. 2, where we present our CoRoT space photometry, spectropolarimetric observations, and the data reduction, while the physical parameters of the target star HD 50773 and a detailed abundance analysis are presented in Sect. 3. In Sect. 4 we discuss our analysis of the surface structure of HD 50773 by applying Bayesian methods to the CoRoT light curve. Details about the magnetic field geometry and the surface abundance structures of individual elements are presented in Sects. 5 and 6. In Sect. 7 we discuss the theoretical predictions of light variability, taking into account the surface abundance inhomogeneities, and Sect. 8 is devoted to the comparison of our different approaches and the discussion of results.

2. Observations and data reduction

2.1. CoRoT Data

HD 50773 was observed by CoRoT from February 3 to April 2, 2007. The reduction of the CoRoT photometry to the N2 level was performed as described in Appourchaux et al. (2008). The time span of 57.7 days covers 27.6 stellar rotational cycles, and the overall variation in brightness amounts to 0.017 mag. The wavelength range (in the CoRoT seismology field) de-

finied by the CoRoT photometric CCD covers 2500Å to 11000Å. For additional instrumental details, we kindly refer the reader to Fridlund et al. (2006).

Any instrumental signal or stellar variation not associated with rotation and not consistent with Gaussian noise or a linear trend would affect the determination of the various parameters in our spot model. Therefore, we tested the light curve for any jumps in flux, which could either be caused by the star itself or the instrument. We first subtracted a linear trend from the time series attributed to aging effects in the detector and electronics chain (Auvergne et al. 2009) and subsequently removed the apparent stellar rotational variability from the data. Low-order polynomials were fitted to short subsets (~ 0.2) of the residual data in order to eliminate the signal changes due to rotation, which resulted in a residual time series shown in the upper panel of Fig. 2. In a next step we calculated a running average of this residual time series using a box width of 370 data points and subtracted the running average - resampled to the original time tags - from the original light curve. The result is shown in the lower panel of Fig. 2.

The upper panel of Fig. 3 shows the original CoRoT light curve with the cumulated corrections obtained so far as a thick black line. After the subtraction of these corrections we arrive at the final CoRoT light curve of HD 50773 (lower panel of Fig. 3). To reduce the large amount of 139271 data points (and the computing time) we binned the data as follows. The un-binned autocorrelation function shows a sharp drop within 0.0006 days and goes through zero at a timelag of 0.02 days. As a compromise between high time resolution - at least a fourth of CoRoT's orbital period, i. e. 0.018 days - and statistical independence, a binning length of ≈ 0.01 days was chosen. Each timebin is represented by the median, and the assigned weight is simply the number of data used, which on average were 26.7 original data points. The binned light curve is used for our analysis described in the following sections.

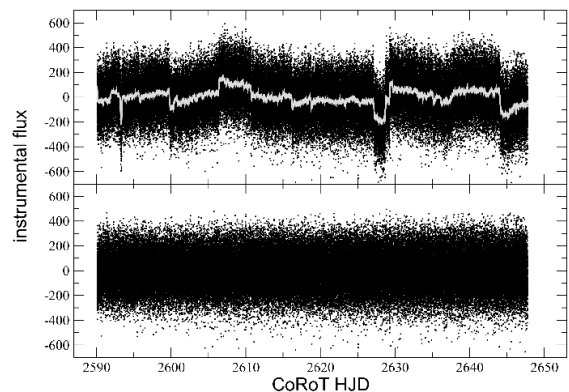


Fig. 2. Upper panel: The CoRoT light curve of HD 50773 after removal of the obvious stellar variability (black). Some small flux changes are remaining, which are shown in grey after applying a running average (see text in this Sect.). Lower panel: After the subtraction of this running average, the residuals to the stellar variability are mostly consistent with white noise. Note that the abscissa in this plot is in CoRoT - HJD (= HJD - 2451545.0) and the ordinate in CoRoT - internal flux units of the N2 data level after subtraction of the mean photometric signal, which is 301700 in the same units.

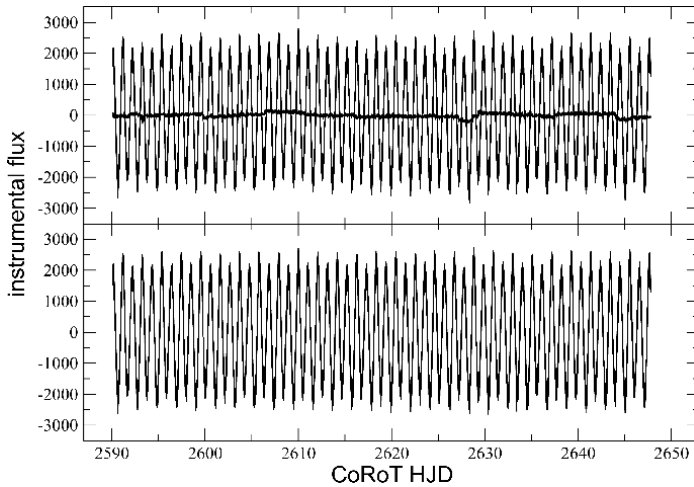


Fig. 3. *Upper panel:* The detrended CoRoT light curve of HD 50773 and the resampled running average containing information about residual signal obviously not related to stellar rotation (thick black line). *Lower panel:* The final light curve containing only signal due to stellar rotation and used for the Bayesian analysis, but zero-mean corrected for the sake of comparability. Again the abscissa in this plot is in CoRoT - HJD (= HJD - 2451545.0) and the ordinate in CoRoT - internal flux units of the N2 data level after subtraction of the mean photometric signal, which is 301700 in the same units.

2.2. Spectropolarimetry with NARVAL, ESPaDoNS and SemelPol

Spectropolarimetric observations of HD 50773 were obtained at the Canada-France-Hawaii-Telescope (CFHT) using ESPaDoNS (Donati et al. 2006) and NARVAL, which is attached to the Telescope Bernard Lyot (TBL) at Pic du Midi (February 2007, simultaneously to the COROT observations), and in December 2007 with NARVAL and SemelPol in combination with UCLES at the Anglo-Australian Telescope (AAT).

ESPaDoNS and NARVAL are twin spectropolarimeters, consisting of a Cassegrain polarimetric module and a fiber-fed échelle spectrometer allowing the whole (polarimetrically analysed) spectrum from 3700Å to 10000Å to be recorded in each exposure. ESPaDoNS and NARVAL were used in polarimetric mode with a spectral resolution of $R \approx 65000$. Stokes I (unpolarised) and Stokes V (circularly polarised) parameters were obtained by means of four sub-exposures between which the retarders (Fresnel rhombs) were rotated in order to exchange the beams in the whole instrument and to reduce spurious polarisation signatures. The extraction of all spectra was done using Libre-ESpRIT (Donati et al. 1997), a fully automatic reduction package installed both at CFHT and TBL.

SemelPol is a visitor instrument, which is mounted at the Cassegrain focus of the AAT in combination with the UCLES spectrograph. This combination has been described in detail by e.g. Semel et al. (1993), or Donati et al. (1999 and 2003). In this setup, two beams of opposite polarisation feed into the échelle spectrograph through two separate optical fibres. During the observing sequence consisting of four subexposures to obtain Stokes I and V, the azimuth of the quarter-wave plate is switched back and forth between $+45^\circ$ (first and fourth exposure) and -45° (exposures 2 and 3), allowing for a removal of systematic errors in the measurements. Spectra cover the wavelength range be-

tween $\approx 4300\text{\AA}$ and 6800\AA with a resolving power of $R \approx 70000$. The above mentioned Libre-ESpRIT (Donati et al. 1997) package was used for data reduction.

We observed HD 50773 during 12 nights and obtained 19 Stokes V (and Stokes I) series (Table 1).

For the Zeeman analysis Least-Squares Deconvolution (LSD, Donati et al. 1997) was applied to each observation. To perform the cross-correlation analysis we produced a mask calculated for physical parameters as deduced in Sect. 3 and abundances obtained in Sect. 3.2, using spectral line lists from the Vienna Atomic Line Database (VALD; Piskunov et al. 1995; Kupka et al. 1999; Ryabchikova et al. 1999). The longitudinal magnetic field (B_l) measurements with their 1σ error bars in G were computed using the first-order moment method (Rees & Semel 1979, Donati et al. 1997).

Rotation phases of HD 50773 (Table 1) were calculated according to the ephemeris and rotation period (derived by us):

$HJD = 2454135.09 + 2^d09101 \times E$, where E is an integer number.

3. Atmospheric parameters and abundance analysis

3.1. Atmospheric parameters

In order to derive accurate atmospheric parameters for HD 50773, a grid of model atmospheres centered on $T_{\text{eff}} = 8500\text{K}$ and $\log g = 4.1$ was computed using LLMODELS (Shulyak et al. 2004). Synthetic spectra based on these models were calculated with synth3 (Kochukhov 2007) and were compared to observations. Atomic parameters used for spectrum synthesis were extracted from VALD using the default configuration file. Average surface element abundances of spectral lines corresponding to different species, which could later be used as starting values for Doppler imaging, were determined from equivalent width measurements and an adapted version of the WIDTH9 code (Kurucz 1993). For this analysis observations from different phases were co-added to reduce line profile asymmetries. The effective temperature was derived from abundance - excitation potential correlations computed for all model atmospheres in the grid using 39 Fe I and 15 Fe II lines. This correlation is very sensitive to changes in effective temperature. For further analysis we adopted a model with $T_{\text{eff}} = 8300\text{K}$, for which the calculated Fe abundance was found to be independent of the excitation potential of individual transitions.

For abundance analyses of normal A type stars the surface gravity $\log g$ is usually determined via the ionisation equilibrium of the Fe lines, i.e., for a correct value of $\log g$ the iron abundance determined from a set of Fe I lines is equal to the abundance derived from Fe II lines. In the atmospheres of Ap stars, diffusion processes lead to a vertical stratification of Fe, Cr and other elements. Spectral lines corresponding to neutral and ionised species of the same element sample different atmospheric layers. Therefore abundances determined from these two different line sets will not be equal, even for a correct choice of $\log g$. For our analysis we decided to use the value of $\log g = 4.1 \pm 0.1$, based on Geneva photometry (Burki et al. 2009, in prep.).

Due to the rather high $v_e \sin i$ of HD 50773, possible vertical abundance stratification within the stellar atmosphere had to be neglected, which leads to a minimum error of $\pm 300\text{K}$ for T_{eff} .

The smallest slope in the equivalent width-abundance correlation was obtained for microturbulent velocities between 2 and 3 km s^{-1} . This rather large value is possibly due to a combination of the effects of line broadening by the magnetic field and the

Table 1. Journal of the spectroscopic observations of HD 50773: dates, instrument, exposure time, peak S/N, HJD, phase, B_1 , σ .

Date	Instrument	Exposure time	S/N	HJD	Phase	B_1	σ
				(245 0000+)		G	G
22 Feb. 2007	NARVAL	4x600s	136	4154.5095	0.289	50	99
28 Feb. 2007	NARVAL	4x800s	78	4159.5142	0.682	-96	223
28 Feb. 2007	ESPaDOnS	4x525s	130	4159.7852	0.812	-148	102
04 Dec. 2007	NARVAL	4x600s	64	4439.4471	0.557	-179	188
05 Dec. 2007	NARVAL	4x600s	81	4439.7077	0.681	-70	161
12 Dec. 2007	NARVAL	4x600s	156	4447.4335	0.376	382	84
13 Dec. 2007	NARVAL	4x600s	154	4447.6877	0.498	441	69
13 Dec. 2007	NARVAL	4x600s	141	4448.4290	0.852	-386	94
14 Dec. 2007	NARVAL	4x600s	123	4448.6841	0.974	-278	95
14 Dec. 2007	NARVAL	4x600s	163	4449.4622	0.346	219	80
15 Dec. 2007	NARVAL	4x600s	115	4449.7114	0.465	579	98
15 Dec. 2007	NARVAL	4x600s	119	4450.4431	0.815	-45	119
16 Dec. 2007	NARVAL	4x600s	107	4450.6943	0.935	-382	114
16 Dec. 2007	NARVAL	4x600s	110	4451.4378	0.291	-100	138
17 Dec. 2007	NARVAL	4x600s	106	4451.6888	0.411	355	126
18 Dec. 2007	NARVAL	4x600s	170	4453.4605	0.258	9	77
19 Dec. 2007	NARVAL	4x600s	138	4453.7095	0.377	273	94
28 Dec. 2007	SemelPol	4x2100s	170	4463.1133	0.875	-368	43
01 Jan. 2008	SemelPol	4x900s	130	4467.0690	0.766	13	64

co-addition of spectra from different rotation phases. Therefore, abundances obtained with $v_{\text{mic}} = 2 \text{ km s}^{-1}$ were used as starting values for our Doppler imaging.

3.2. Abundance analysis

Using the WIDTH9 code and the co-added spectrum of HD 50773, a crude abundance analysis of 17 species was performed. Due to the strong rotational variability of practically all spectral lines in this star, detailed abundances for individual rotation phases can only be derived via Doppler imaging. The values given in this section can therefore be taken only as first order approximations. In Table 2 we present abundances of 17 species derived from equivalent width measurements in co-added spectra of HD 50773, based on $T_{\text{eff}} = 8300 \text{ K}$, $\log g = 4.1$, and $v_{\text{mic}} = 0 \text{ km s}^{-1}$. For each species we give the number of lines measured and the standard deviations of the derived element abundances. A complete list of lines used for this analysis is given online in Table 3.

4. Analysing the light curve with Bayesian methods – Photometric Imaging (PI)

Fitting the light curve in terms of a few circular spots needs the estimation of many parameters. The star is described by the inclination angle i and the rotational period P . Each spot needs at least three further parameters: longitude λ , latitude β , and spot radius γ (in the following longitude increases in the direction of stellar rotation, and the zero point is the central meridian facing the observer at the beginning of the time series). Additional parameters are the spot intensity κ and the coefficient u in the linear limb-darkening law. We use $u = 0.4415$ derived within model atmosphere calculations with the ATLAS9 code (Kurucz 1993) applying the wavelength range from 2500 \AA to 11000 \AA , which is defined by the CoRoT photometric CCD.

Before analysing the data proper prior distributions have to be assigned. If information on the star's inclination i is missing, $\cos i$ is evenly distributed for $0 \leq i \leq \pi/2$. Concerning a latitude β , the value of $\sin \beta$ is assumed to be evenly distributed. All non-dimensionless parameters like radii or frequencies are

Table 2. Abundances of 17 species derived from equivalent width measurements in co-added spectra of HD 50773.

Species	$\log N_{\text{elm}}/N_{\text{tot}}$	# of lines	stdev
C I	-3.03	1	-
Mg I	-4.49	2	0.04
Si I	-4.14	2	0.22
Si II	-3.93	2	0.18
Ca I	-5.60	1	-
Sc II	-8.32	1	-
Ti II	-6.70	2	0.21
V II	-6.74	1	-
Cr I	-4.49	2	0.33
Cr II	-5.10	7	0.32
Mn I	-6.36	1	-
Fe I	-4.00	39	0.30
Fe II	-3.91	15	0.37
Ni I	-5.48	1	-
Y II	-9.13	1	-
Pr III	-6.57	1	-
Nd III	-7.82	1	-

represented by their logarithms. This makes certain that the posterior distribution for a radius will be consistent with that of an area and likewise the posterior for a frequency with that of a period, i. e. it does not make a difference whether one prefers radii or areas, frequencies or periods.

The likelihood function is constructed as follows. Spotted stars which are geometrically similar exhibit the same light curve, except for an offset in magnitude. This offset is considered irrelevant and removed by integration. This is even indicated, as the magnitude of the unspotted star, the zero point, is unknown. This integration can be done analytically if the measurement errors are assumed to be Gaussian-distributed over the magnitudes.

With the N data points d_i gained at times t_i , their standard deviations σ_i , the fit $f_0(t_i)$ and an offset c_0 , the likelihood is represented by:

$$\Lambda(\sigma_{1..N}, c_0, p_1 \dots p_M; d_i) =$$

$$\prod_{i=1}^N \frac{1}{\sqrt{2\pi}\sigma_i} \exp\left(-\frac{(d_i - f_0(t_i, p_j) - c_0)^2}{2\sigma_i^2}\right).$$

The unknown parameters are denoted by p_j , with $j = 1 \dots M$, with $M = 24$ in the case of four spots. We set $\sigma_i = s_i \cdot \sigma$, with the relative errors s_i being normalised according to $\sum_{i=1}^N 1/s_i^2 = N$.

First integrating analytically the measurement error σ – using Jeffreys’ $1/\sigma$ -prior – and then the uninteresting offset c_0 , one gets a likelihood depending only on the spot modelling parameters $p_1 \dots p_M$. It is this *mean* likelihood, averaged appropriately over measurement error and offset, from which the posterior density distributions for all unknowns is obtained by marginalization. This is a great advantage of the Bayesian approach: It provides not only a set of most probable parameter values but moreover *expectation* values and error bars exclusively from the data. We can argue that the method itself determines error and offset from the data.

The Markov chain Monte Carlo (MCMC) method (cf. Press et al. 2007) has been applied to explore the likelihood mountain in a high-dimensional parameter space. The MCMC technique has already shown its capabilities when analysing photometric data from the Canadian MOST satellite (Croll 2006, Fröhlich 2007).

A set of 64 Markov chains was generated. Each chain has to perform some 10^7 steps, and after a burn-in period every thousandth successful step was recorded in order to suppress the correlation between successive steps. Budding’s star-spot model (1977) was used to model the light curve.

4.1. Results from Bayesian PI

There are two solutions which fit the light curve equally well: one with three “dark” spots and another one with four “bright” spots. In both cases the residuals amount to 0.12 mmag. The autocorrelation function of the residuals reveals that there is even more in the data than what can be represented by a simple model with circular spots. From a formal point of view the solution with “dark” spots is the more probable one, because the fit needs only three spots instead of four. But in view of the spectroscopic evidence the four-spot solution makes much more sense.

The expectation values with $1-\sigma$ confidence limits for the star and the starspot parameters are presented in Table 4. From spot longitudes, spot epochs (HJD) can be computed:

$$E_i = 2454135.09 - 2.09101 \cdot \lambda_i / 360^\circ.$$

The reader should be aware that the estimated parameter values and their (surprisingly small) error bars are those of the model constrained by the data. They make sense *given the model is true*, i.e. that there are four circular spots, a linear limb-darkening law with prescribed coefficient holds and so forth.

In a further run the coefficient u of the limb-darkening law has been determined from the data itself. This does not lead to any improvement of the fit. The fact that the deduced limb-darkening coefficient $u = 0.405^{+0.010}_{-0.013}$ proves to be only a little bit smaller than the theoretical value of 0.4415 strengthens the confidence in the reliability of the four-spot model.

The high accuracy of the photometric data makes it reasonable to question the common wisdom and look for spot evolution as well as differential rotation and even for a period drift in an Ap star. Time evolution is described by a Legendre expansion of the logarithm of the spot areas up to the seventh power in time. We have done this in the case of the three dark spots only. The result is a clear null result: we neither find a spot area evolution

Table 4. Spot parameters: we list *expectation* values and $1-\sigma$ uncertainties, the period P is given in days, the spot intensity is presented in units of the intensity of the unspotted surface.

parameter		expectation and uncertainties	
inclination	i	40:3	+1.3 -0.2
1st longitude	λ_1	184:6	+0.4 -0.4
2nd longitude	λ_2	140:2	+0.5 -0.5
3rd longitude	λ_3	32:7	+0.3 -0.3
4th longitude	λ_4	312:8	+0.3 -0.3
1st latitude	β_1	18:5	+0.8 -1.0
2nd latitude	β_2	21:3	+1.0 -1.0
3rd latitude	β_3	-7:3	+0.4 -0.4
4th latitude	β_4	-20:2	+1.0 -1.1
1st radius	γ_1	23:9	+0.5 -0.4
2nd radius	γ_2	15:0	+0.5 -0.5
3rd radius	γ_3	31:3	+0.4 -0.4
4th radius	γ_4	39:3	+0.8 -0.8
period	P	2.091007	+0.000007 -0.000007
spot intensity	κ	1.22	+0.01 -0.02

Table 5. CoRoT orbit effect. The estimated CoRoT orbital period is given in min, amplitudes in ppm.

parameter		expectation and uncertainties
CoRoT orbital period	P_{orbit}	103.0660±0.0012
amplitude (P_{orbit})	A_1	111±2
amplitude ($P_{\text{orbit}}/2$)	A_2	33±2
amplitude ($P_{\text{orbit}}/3$)	A_3	8 $^{+3}_{-2}$
amplitude ($P_{\text{orbit}}/4$)	A_4	13±2

above a 2.5- σ level nor a hint for any departure from rigid rotation. The spot periods are constant at $\lesssim 3$ ppm. The lapping time for the two largest (“dark”) spots would have to exceed 100 years.

It is important to clean the data from a CoRoT orbit effect: The residuals drop by a factor of five by subtracting a periodic signal containing CoRoT’s orbital period as well as three overtones (Table 5).

5. Magnetic field geometry

All 19 Stokes V LSD-line-profiles derived from ESPaDOnS, NARVAL, and SemelPol spectropolarimetric observations are grouped in a single time-series, in order to increase the rotational sampling as much as possible and build a proper data set for tomographic inversion. The time-series is then used to reconstruct the surface magnetic geometry of the star by means of Magnetic Doppler imaging (Donati & Brown 1997).

To compute synthetic Stokes V line profiles, we make a model of a spherical stellar surface divided into a grid of pixels, each pixel producing a local Stokes I and V profile. Assuming a given magnetic field strength and orientation for each pixel, local Stokes V profiles are calculated under the assumption, valid in the weak field limit, that Stokes V is proportional to $g \cdot \lambda_0^2 \cdot B_{\parallel} \cdot \partial I / \partial \lambda$, where λ_0 is the average wavelength of the LSD profile (about 534 nm for HD 50773), B_{\parallel} is the line-of-sight projection of the local magnetic field vector, g is the effective Landé factor of the LSD profile (equal to 1.2) and $\partial I / \partial \lambda$ the wavelength derivative of the local Stokes I line profile. We further assume that there are no large-scale brightness or abundance inhomogeneities over the stellar surface, so that synthetic

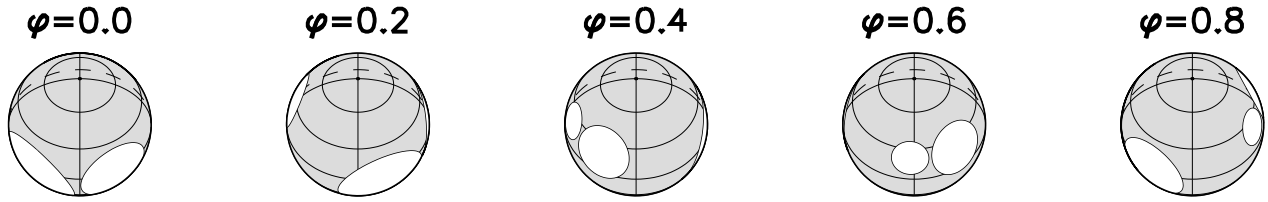


Fig. 4. Locations of the four bright photometric spots, assumed to be of circular shape, plotted at five equidistant phases.

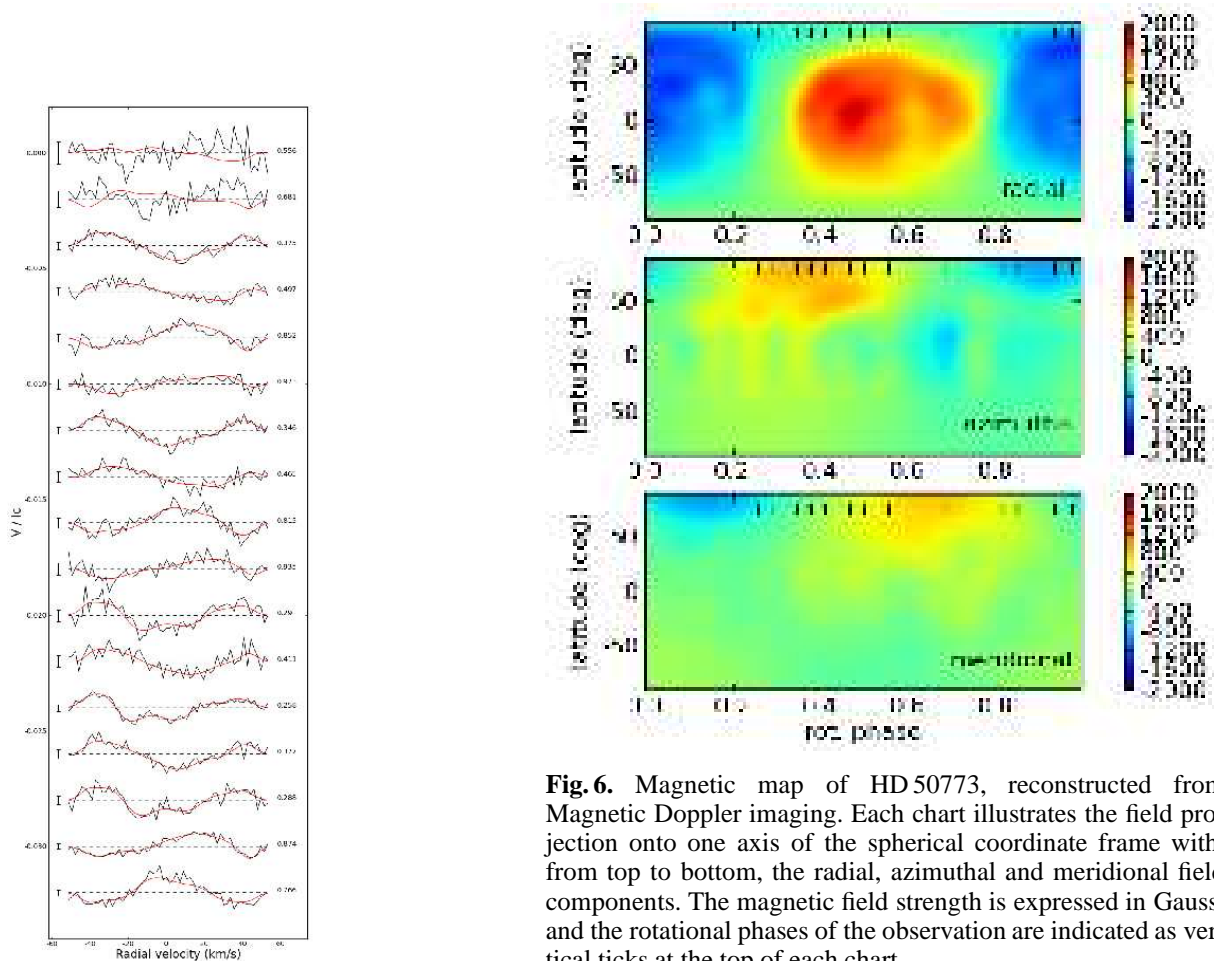


Fig. 5. Stokes V profiles of HD 50773 after the correction of the mean radial velocity of the star. Black lines represent the data and red/grey lines correspond to the synthetic profiles of our magnetic model. Successive profiles are shifted vertically for better visibility. Rotation phases of observations are indicated in the right part of the plot, and error bars are illustrated on the left of each profile.

Stokes I profiles are locally the same over the whole photosphere. However we would like to mention that some features in the Stokes V profile variability may arise from the influence of chemical spots and not from the magnetic topology itself, which has been demonstrated in Kochukhov et al. (2002).

Fig. 6. Magnetic map of HD 50773, reconstructed from Magnetic Doppler imaging. Each chart illustrates the field projection onto one axis of the spherical coordinate frame with, from top to bottom, the radial, azimuthal and meridional components. The magnetic field strength is expressed in Gauss, and the rotational phases of the observation are indicated as vertical ticks at the top of each chart.

For each pixel, we impose a Gaussian shape to the local synthetic Stokes I line profile. Parameters of the local profile (width and depth) are chosen to optimise the adjustment of Stokes I LSD profiles of HD 50773. The amplitude of the local Stokes profiles for pixels located at the visible hemisphere of the star is then weighted according to a linear limb-darkening coefficient equal to 0.49, and their central wavelengths are shifted according to the line-of-sight velocities of individual pixels, assuming $v \sin i = 46.3 \text{ km s}^{-1}$ (a value minimising the χ^2 of the magnetic model) and $i = 40^\circ$.

Synthetic Stokes V profiles are computed for the 19 observed rotation phases and are compared to the observations. The data adjustment is iterative and based on a maximum entropy algorithm (Skilling & Bryan 1984). The version of the code used here makes a projection of the surface magnetic field onto a spherical harmonics frame (Donati et al. 2006), with the magnetic field

geometry splitted between a poloidal and toroidal component (Chandrasekhar 1961). We limit the spherical harmonics expansion to $\ell < 10$. Using this procedure, the spectropolarimetric data are adjusted at a reduced $\chi^2 = 1.2$ (Fig. 5). The final χ^2 , as well as the reconstructed magnetic topology, are left essentially unchanged as long as the adopted spherical harmonics description respects $\ell_{\max} \geq 6$. The reconstructed magnetic geometry is illustrated in Fig. 6, from which we can easily see that the surface magnetic field of HD 50773 is mostly a highly inclined dipole, with a polar strength of about 2 kG.

A closer look at the weight of the complex spherical harmonics coefficients $\alpha_{\ell,m}$, $\beta_{\ell,m}$, and $\gamma_{\ell,m}$ (defined by Donati et al. 2006) is however necessary to obtain a more quantitative information on the magnetic field distribution. Whereas the inversion procedure allows for the reconstruction of both poloidal and toroidal components of the field, 96% of the reconstructed magnetic energy ends up in the poloidal component. As expected for such a marginal amount of toroidal field, the outcome of the inversion code is actually almost similar if we impose the more restrictive case of a purely poloidal field. As can be readily seen on the map, most of the magnetic energy (91 %) is stored in the dipolar component. Higher-order field components are however necessary to obtain a convincing modelling of the observations (up to $\ell = 6$). Finally, the predominance of a highly non-axisymmetric field distribution translates into a low amount of magnetic energy in axis-symmetric spherical harmonics modes ($m = 0$), which do not gather more than 2% of the overall photospheric magnetic energy.

6. Spectroscopic Doppler imaging

By applying the *Doppler imaging* technique (DI), we are able to translate the partially very pronounced variations in the spectral line profiles of HD 50773, linked to stellar rotation, into surface maps of the abundance distribution. The longitude of a spot is directly deduced from the wavelength position of the distortion within the profile, whereas its latitude can only be derived from time-series observations. In the case of HD 50773, 19 such spectra (listed in Table 1) could be used for the inversions with INVERS12, the DI code we used (Kochukhov et al. 2004b). In this code, where specific intensities are calculated for each visible surface element for each iteration, it is possible to simultaneously calculate abundance maps of several chemical elements even from blended spectral lines. Mapping HD 50773, we could derive surface abundance distributions for Mg, Si, Ca, Ti, Cr, Fe, Ni, Y and Cu. We have to mention that Ti and Ni were derived from spectral line blends where these elements contribute only moderately to the overall absorption. Input parameters were determined as described in Sect. 3, $v_e \sin i$ and inclination i used as in Sect. 5.

6.1. Surface abundance structures of individual elements

6.1.1. Magnesium and calcium

The surface abundance distribution of *magnesium* was modeled using a blend of Mg I with Fe I and Y II around 5400 Å. We find a variation between $\log N_{\text{Mg}}/N_{\text{tot}} = -5.2$ and -3.3 dex and the element, contrary to the other mapped species, does not exhibit a clear correlation with the magnetic field geometry, and is enhanced in a spotted, belt-like structure near the rotational equator (Fig. 6.1). The abundance distribution of *calcium*, variable between $\log N_{\text{Ca}}/N_{\text{tot}} = -6.6$ and -2.7 dex and determined from Ca I and II blended with Fe I around 6102 Å and 6456 Å also

shows a belt-like structure at the rotational equator with a pronounced region of overabundance at the positive magnetic pole.

6.1.2. Iron-peak elements

A blend of *titanium* II with Fe I and Cr II at 5169 Å was used to recover the surface abundance structure of this element. Ti II shows a tendency to accumulate at the magnetic poles and to avoid the magnetic equatorial regions, varying between $\log N_{\text{Ti}}/N_{\text{tot}} = -8.3$ and -6.1 dex.

Chromium was mapped using several single Cr lines at 5237 Å and 5280 Å plus a blend of Cr II with Fe I, Y II and Ni II around 5510 Å and was found to be variable between $\log N_{\text{Cr}}/N_{\text{tot}} = -6.8$ and -3.0 dex. The abundance distribution of the element shows as that of Fe a clear correlation to the magnetic field geometry, being enhanced at the poles and depleted at the magnetic equator, whereby the two enhancement regions seem to be subdivided into two spots.

The surface abundance of *iron* was determined using single Fe I lines at 5383 Å and 5400 Å and again the blend of Fe I with Cr II, Y II, and Ni II at 5510 Å. As mentioned above, Fe is, as Cr, closely correlated with the magnetic field, avoiding the magnetic equatorial region. The elemental abundance varies between $\log N_{\text{Fe}}/N_{\text{tot}} = -5.3$ and -1.9 dex.

Nickel, derived from the blend of this element with Cr II, Fe I, and Y II at 5510 Å and variable between $\log N_{\text{Ni}}/N_{\text{tot}} = -6.1$ and -5.6 dex shows an opposite correlation to the magnetic field geometry from what we saw from Cr and Fe: it is enhanced at the magnetic equatorial regions and depleted where the positive and the negative magnetic pole cross the line of sight.

6.1.3. Yttrium and copper

A similar distribution as that for Ni can be observed for *yttrium*. Varying from $\log N_{\text{Y}}/N_{\text{tot}} = -9.9$ to -7.7 dex it also shows a pronounced region of overabundance at the magnetic equator and depletion on the poles. The element was mapped from a blend with Cr II, Fe I, and Ni II at 5510 Å and 5662 Å. *Copper* around 5153 Å, blended with Cr II, Ti II, and Fe I was used to derive the surface abundance of this element, and we see it to vary between $\log N_{\text{Cu}}/N_{\text{tot}} = -8.0$ and -6.9 dex. At first glance, it seems to be closely related to the distributions of Ni and Y, being enhanced at the magnetic equator, but it exhibits in addition two less pronounced spots close to the magnetic poles.

7. Effect of abundance inhomogeneities on spectral energy distribution

As already mentioned in Sect. 1 the light curve of HD 50773 obtained by CoRoT has a periodic form with two clear maxima of slightly different amplitudes at phases $\phi \approx 0.05$ and 0.52 (see Fig. 1). This photometric variability is likely to be connected with inhomogeneous surface element distributions, similar to those recently reported and successfully modeled by Krtićka et al. (2007) for the hot CP2 star HD 37776. The physical nature of this effect is directly connected with the radiative flux redistribution due to enhanced or deficient opacity in the abundance spots relative to the rest of the stellar surface. Hence, as a star rotates, the observer sees different stellar regions that are emitting a different amount of radiative flux producing characteristic variability of indices in phase-resolved photometry.

Of course, abundance spots are not the only effect which may produce visible rotational modulation in photometry. For

Table 6. Table of elements and spectral lines used for mapping of HD 50773: element, spectral lines used, the $\log gf$ values of these lines, abundance interval in dex, for comparison the solar values (Asplund et al. 2005) are presented. Atomic parameters used in our study were extracted from VALD.

Species	λ (Å)	$\log gf$	E_{low} (eV)	$\log(N/N_{\text{tot}})$ min, max	blended with	$\log(N/N_{\text{tot}})$
Mg I	5401.5210	-0.340	11.6300		Fe, Y	-4.51
Si II	5055.9840	0.593	10.0740	-4.8, -2.6		-4.53
Si I	5421.1680	-2.250	5.6190		Cr, Mn	
Si I	5421.3830	-1.480	5.6190		Cr, Mn	
Ca I	6102.7230	-0.862	1.8790	-6.5, -2.7	Fe	-5.73
Ca II	6456.8750	-0.539	8.4380		Fe	
Ti II	5154.0680	-1.750	1.5660	-8.3, -6.1	Cr, Fe	-7.14
Cr II	5237.3290	-1.350	4.0730	-7.1, -3.6		-6.40
	5280.0540	-2.316	4.0740			
	5510.7020	-2.614	3.8270		Fe, Y, Ni	
Fe I	5383.3692	0.645	4.3120	5.4, -3.6		-4.59
	5400.5022	-0.160	4.3710			
	5400.6560	-2.482	3.6350			
	5401.2689	-1.920	4.3200			
Ni I	5510.0030	-0.900	3.8470	-6.1, -5.6	Cr, Fe, Y	-5.81
Y II	5509.8950	-1.010	0.9920	-9.9, -7.5	Cr, Fe, Ni	-9.83
	5662.9250	0.160	1.9440		Fe	
Cu II	5153.2300	0.217	3.7860	-8.0, -6.9	Cr, Ti, Fe	-7.83

instance, strong magnetic fields may also influence the total radiative flux gradients along the stellar surface since the magnetic opacity and flux distribution strongly depend upon the modulus of the magnetic field (Kochukhov et al. 2005, Khan & Shulyak 2006). However, the mean surface magnetic field of HD 50773 is too weak to noticeably affect the integrated opacity coefficient, and thus the chemical spots are the most probable source of the observed variability. Details of our modelling of the light variability taking into account surface abundance variations of HD 50773 are presented in this Sect..

Having produced abundance maps using the Doppler imaging technique it is, in principle, possible to directly model the light curve variability in the same manner as it was done in Krtićka et al. (2007). However, in case of HD 50773, there are several elements that play a noticeable role in the flux redistribution, which would require the calculation of hundreds of model atmospheres for an accurate surface integration. These extensive computations are out of the scope of the present Paper. Nevertheless, to explore the role of the mapped elements on the total energy balance in the atmosphere of HD 50773 (and thus on the ability of these elements to affect the radiative balance) we computed a set of model atmospheres with different assumptions about abundance patterns. To carry out such calculations we employed the LLMODELS stellar model atmosphere code (Shulyak et al. 2004), which incorporates treatment of individual abundance patterns and computes the line opacity in a fine frequency grid for better integration of the radiation field quantities.

First we computed the model with the surface averaged abundances of all nine elements used in our DI analysis. Then another seven models were computed, individually increasing every element by +2 dex relative to its mean value. Once these computations were done, we integrated the output radiative fluxes from all the models and then compared them to the flux produced by the mean abundance reference model. The flux integration was performed in the wavelength range defined by the CoRoT photometric CCD, i.e. from 2500Å to 11000Å, with the resolution of 0.1Å, which is the default flux resolution in

LLMODELS. The effective temperature of all the models was kept to be the same.

We find that among the nine elements considered in this investigation only four have a clear influence on the model energy distribution by more than 1%. These elements are Cr (1.16%), Fe (5.63%), Mg (3.02%), and Si (2.21%). The influence of the other five elements is much less pronounced like, e.g. Ca (0.22%). This agrees well with the recent study of Khan & Shulyak (2007) who investigated the effect of individual chemistry on properties of CP stars and showed the relative importance of Fe, Cr, and Si opacity for all the models considered.

As an example, Fig. 9 demonstrates the synthetic energy distribution computed with enhanced abundances of Cr, Fe, Mg, and Si. In particular, one can note the strong impact of Fe and Si on the shape of the energy distribution. However, the relatively small changes in integrated flux values listed above are due to the fact that the CoRoT photometry covers a wide wavelength region where the flux changes due to enhanced abundance models which act in opposite directions in the UV and the visual regime. We do not show fluxes produced by the other mapped elements since they lie extremely close to the flux of the mean abundance model (thick full line).

8. Conclusions

The high-quality CoRoT data enabled us by applying the Bayesian data analysis to derive stellar surface structures from space quality photometry. Analysing an extensive set of ground based spectropolarimetric data via Doppler imaging and Magnetic Doppler imaging made it possible to directly correlate the results from photometry to the chemical and magnetic stellar surface structure. The resulting abundances were used to model the light variability of HD 50773.

The two areas covered by bright spots found with Bayesian data analysis of the light curve coincide very well with the magnetic polar regions on the surface of HD 50773.

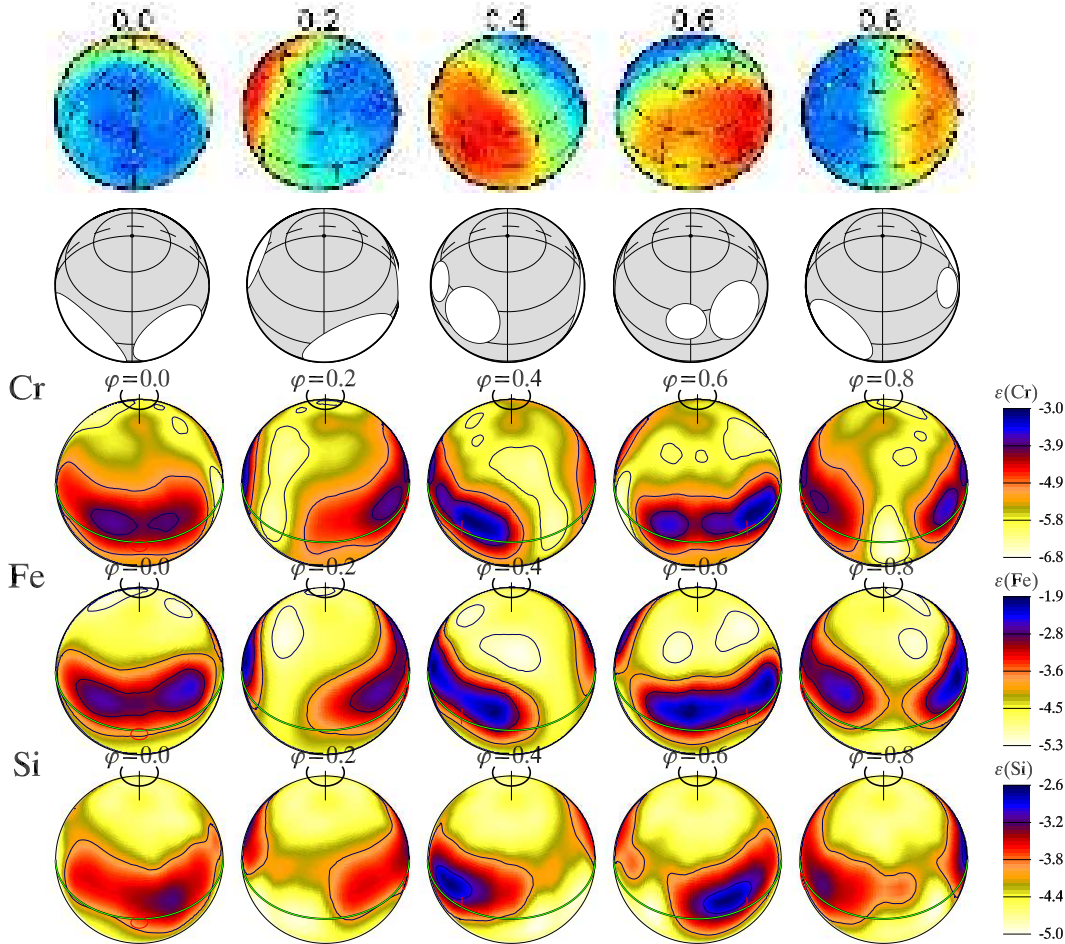


Fig. 7. Top panel: Radial field component of the magnetic map of HD 50773 (as described in Sect. 5). Second panel: locations of the four bright (displayed in black) photometric spots, assumed to be of circular shape. Next three panels: Abundance distribution of Cr, Fe, and Si at the surface of HD 50773 obtained from the lines listed in Table 6 (online). We show the star at an inclination $i = 40^\circ$. Darker areas in the plots correspond to higher elemental abundances, the corresponding scale is given to the right of each panel, and the contours of equal abundance are plotted with steps of 1.0 dex. The circle and the cross indicate the position of the negative and the positive magnetic pole, respectively. All projections are plotted at five equidistant rotation phases.

The elemental abundance spots of the species contributing dominantly to the changes in the integrated flux (Cr, Fe, and Si) are also clearly correlated with the magnetic field geometry and hence with photometric spots. Regions of overabundance are found close to the magnetic poles and those of depletion linked to the magnetic equator. As described in Sect. 7, these elements produce bright spots in the photosphere.

Our results confirm the high potential of combining high quality photometric data obtained in space with ground based spectroscopy. Further studies with a similar approach are currently ongoing and will significantly increase the sample of stars analysed in this way, which is important for modelling elemental diffusion in stars in the presence of a magnetic field.

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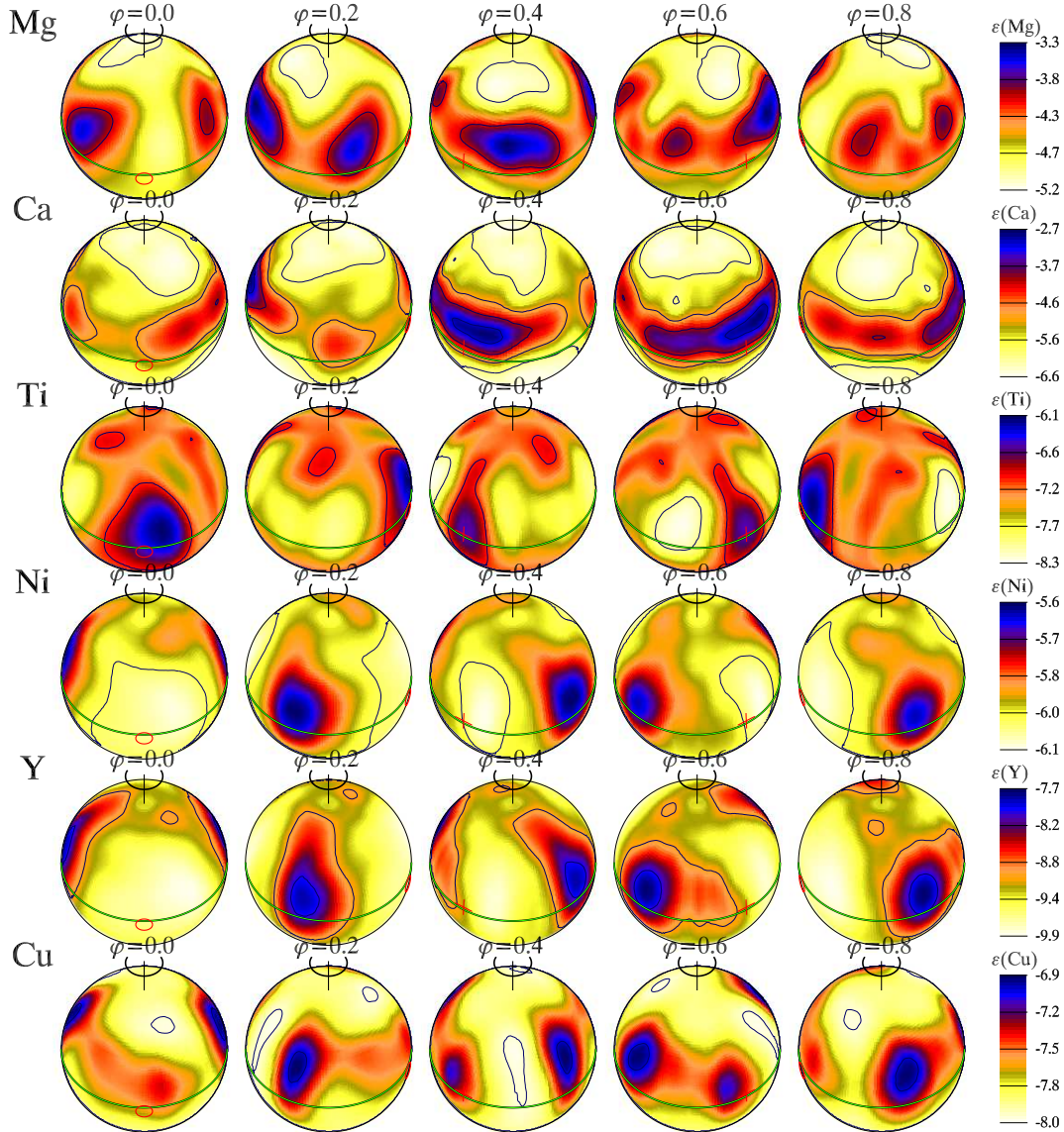


Fig. 8. Same as the three bottom panels in Fig. 7 but for Mg, Ca, Ti, Ni, Y, and Cu.

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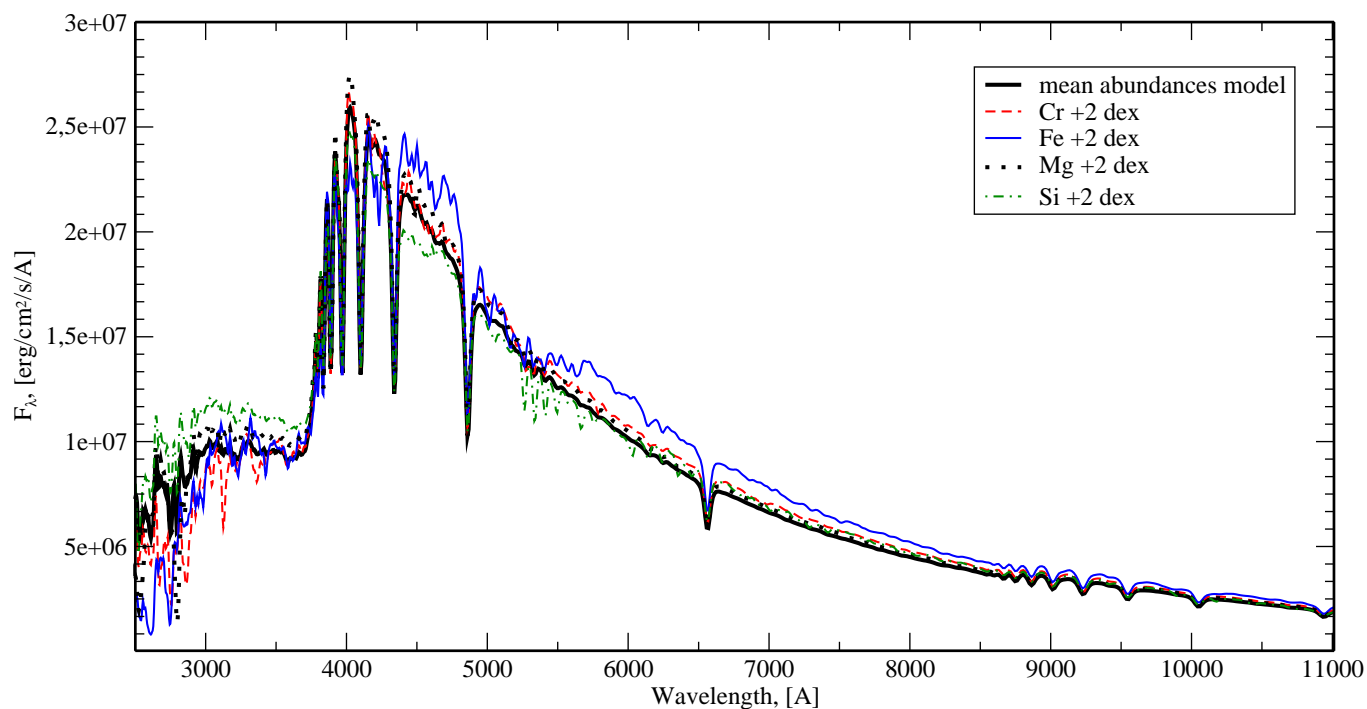


Fig. 9. The synthetic energy distributions of HD 50773 calculated with mean surface abundances and abundances of Cr, Fe, Mg, and Si enhanced by 2 dex. For all models $T_{\text{eff}} = 8300$ K and $\log g = 4.1$ were assumed. To provide a better view all fluxes were convolved with the resolution of $R = 200$.

Online Material

Table 3. Line list used for abundance analysis based on equivalent width measurements in the co-added spectrum of HD 50773. Atomic parameters are taken from the VALD database. For each line wavelength, excitation potential, and $\log gf$ are listed.

Species	λ [Å]	E_{low} [eV]	$\log gf$	species	λ [Å]	E_{low} [eV]	$\log gf$
C I	6001.1179	8.6430	-2.061	Fe I	5445.0424	4.3860	-0.020
				Fe I	5446.9168	0.9900	-1.914
Mg I	5183.6040	2.7170	-0.180	Fe I	5487.7450	4.3200	-0.317
Mg I	5528.4050	4.3460	-0.620	Fe I	5554.8951	4.5480	-0.440
				Fe I	5565.7040	4.6080	-0.213
Si I	5645.6130	4.9300	-1.524	Fe I	5569.6181	3.4170	-0.486
Si I	5772.1460	5.0820	-1.358	Fe I	5572.8424	3.3960	-0.275
Si I	5948.5410	5.0820	-0.780	Fe I	5615.6439	3.3320	0.050
Si II	5055.9840	10.0740	0.593	Fe I	5624.5422	3.4170	-0.755
Si II	5978.9300	10.0740	0.004	Fe I	5633.9465	4.9910	-0.270
				Fe I	5686.5302	4.5480	-0.446
Ca I	5581.9650	2.5230	-0.555	Fe I	5705.9922	4.6070	-0.530
Ca I	5588.7490	2.5260	0.358	Fe I	5762.9922	4.2090	-0.450
Ca I	5594.4620	2.5230	0.097	Fe I	5816.3735	4.5480	-0.601
Ca I	5857.4510	2.9330	0.240	Fe I	5848.1270	4.6080	-1.056
				Fe I	5859.5860	4.5490	-0.419
Sc II	5667.1490	1.5000	-1.309	Fe I	5862.3570	4.5490	-0.127
				Fe I	5914.2010	4.6080	-0.131
Ti II	5185.9018	1.8930	-1.490	Fe I	5930.1799	4.6520	-0.230
Ti II	5418.7675	1.5820	-2.000	Fe I	5934.6549	3.9280	-1.170
				Fe I	6003.0123	3.8810	-1.120
V II	5819.9350	2.5220	-1.703	Fe I	6020.1692	4.6070	-0.270
				Fe I	6024.0580	4.5480	-0.120
Cr I	5206.0370	0.9410	0.019	Fe I	6027.0509	4.0760	-1.089
Cr I	5208.4250	0.9410	0.158	Fe I	6056.0047	4.7330	-0.460
Cr II	5237.3290	4.0730	-1.350	Fe II	5004.1950	10.2730	0.504
Cr II	5310.6870	4.0720	-2.280	Fe II	5061.7180	10.3080	0.284
Cr II	5334.8690	4.0720	-1.826	Fe II	5169.0330	2.8910	-1.303
Cr II	5407.6040	3.8270	-2.151	Fe II	5234.6250	3.2210	-2.230
Cr II	5420.9220	3.7580	-2.458	Fe II	5254.9290	3.2300	-3.336
Cr II	5508.6060	4.1560	-2.252	Fe II	5316.6150	3.1530	-1.850
Cr II	5620.6310	6.4870	-1.395	Fe II	5362.8690	3.1990	-2.616
				Fe II	5427.8260	6.7240	-1.581
Mn I	6021.8190	3.0750	0.034	Fe II	5432.9670	3.2670	-3.527
				Fe II	5439.7070	6.7290	-2.382
Fe I	5065.0180	4.2560	0.005	Fe II	5534.8470	3.2450	-2.730
Fe I	5078.9750	4.3010	-0.292	Fe II	5835.4920	5.9110	-2.702
Fe I	5090.7740	4.2560	-0.400	Fe II	5952.5100	5.9560	-2.388
Fe I	5133.6885	4.1780	0.140	Fe II	5961.7050	10.6780	0.675
Fe I	5139.4628	2.9400	-0.509	Fe II	5991.3760	3.1530	-3.540
Fe I	5162.2729	4.1780	0.020				
Fe I	5202.3360	2.1760	-1.838	Ni I	5715.0660	4.0880	-0.352
Fe I	5232.9403	2.9400	-0.058				
Fe I	5281.7904	3.0380	-0.834	Y II	5087.4160	1.0840	-0.170
Fe I	5353.3736	4.1030	-0.840				
Fe I	5367.4668	4.4150	0.443	Pr III	6053.0044	0.0000	-1.983
Fe I	5383.3692	4.3120	0.645				
Fe I	5393.1676	3.2410	-0.715	Nd III	5845.0201	0.6310	-1.180
Fe I	5434.5238	1.0110	-2.122				